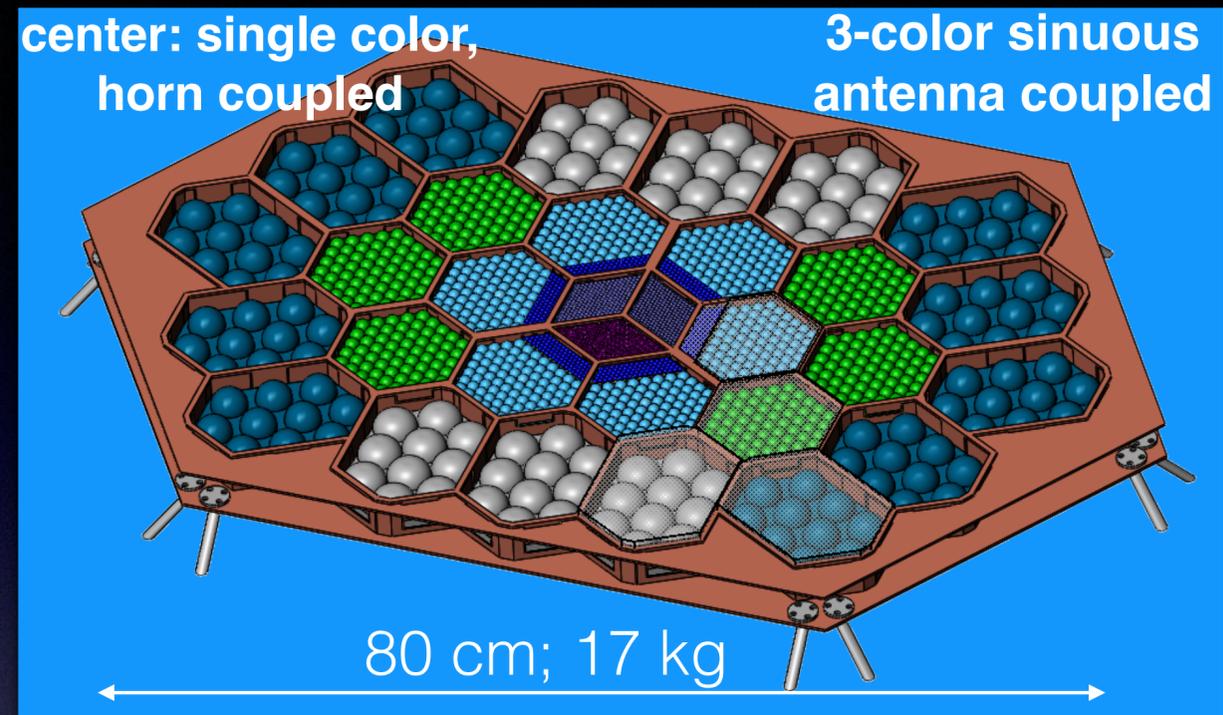


PICO: mm/submm All Sky Imaging Polarimetric Survey

- PICO will produce the deepest maps of Stokes I, Q, U in 21 frequency bands between 20 and 800 GHz
- Maps will have resolution between 38' and 1'.
8 maps, >200 GHz: highest resolution, full sky maps
- Ten redundant surveys: stringent control of systematic errors
- 13,000 transition edge sensor bolometers
- 5 year survey from L2
- Noise baseline: 3300 *Planck* missions ($0.87 \text{ uK} \cdot \text{arcmin}$)
- Noise Current estimate: 6400 *Planck* missions ($0.61 \text{ uK} \cdot \text{arcmin}$)



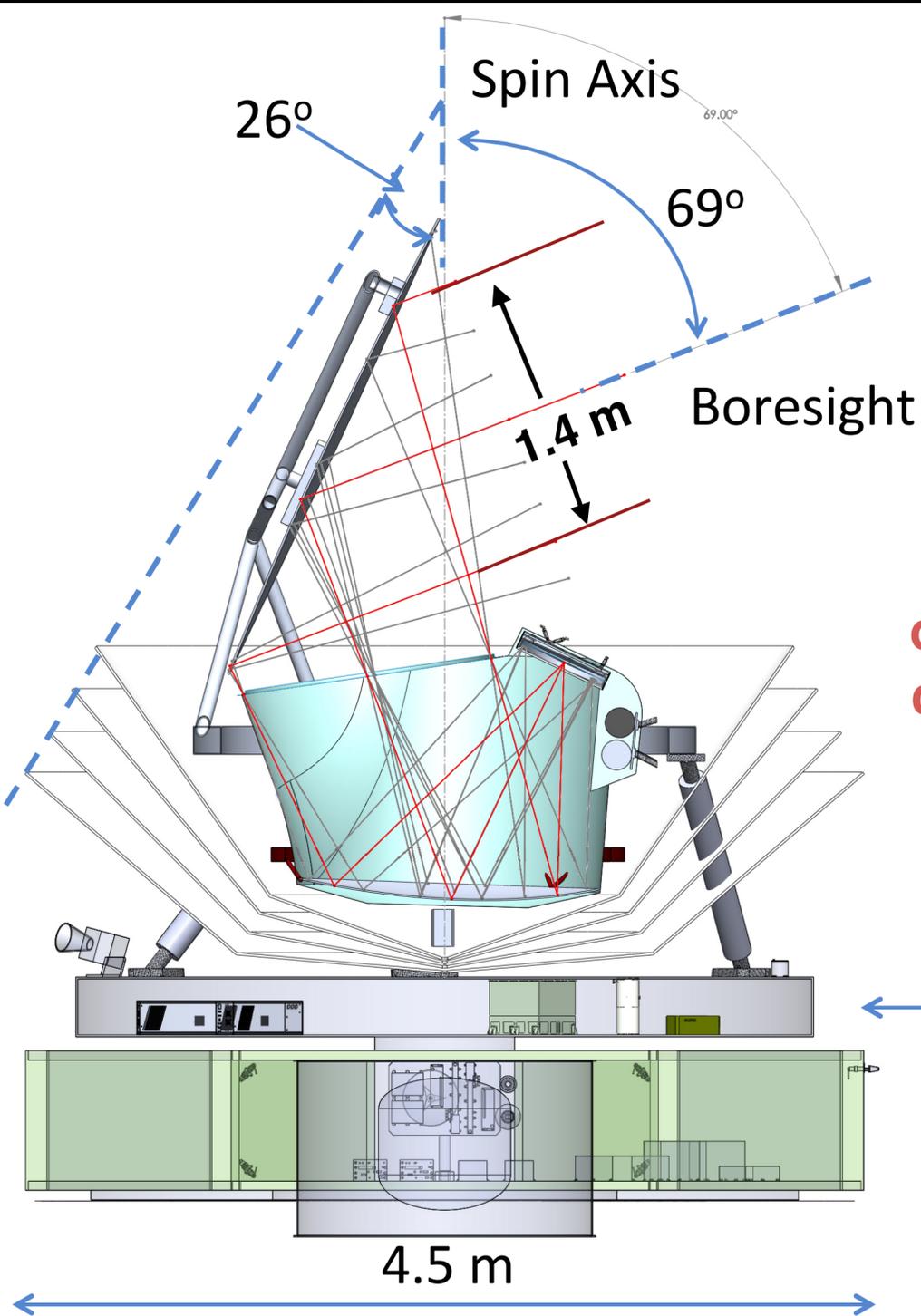
PICO Implementation: Heritage of Planck



Spinning

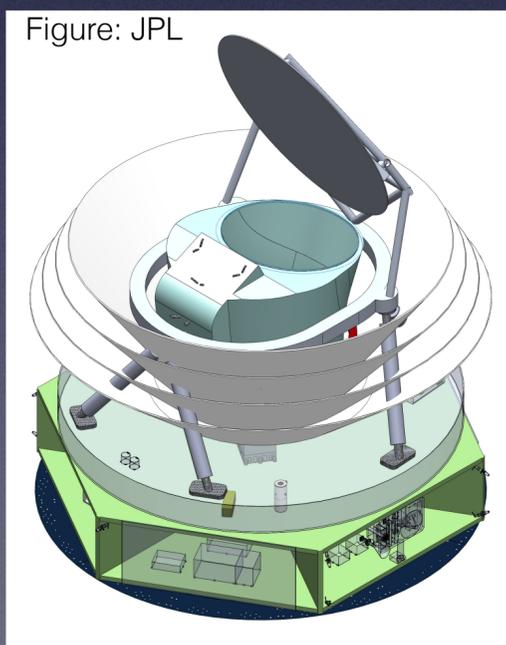
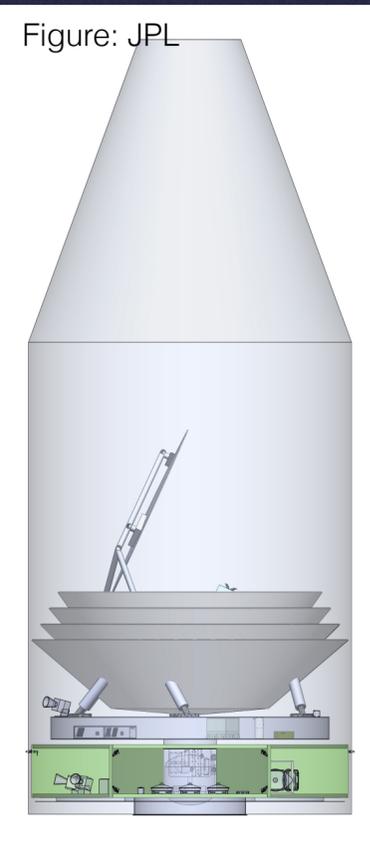
- 2-reflector "Open Dragone" Telescope
- Ambient temperature primary
- 4 K aperture stop
- 4 K secondary reflector
- 0.1 K focal plane (cADR)

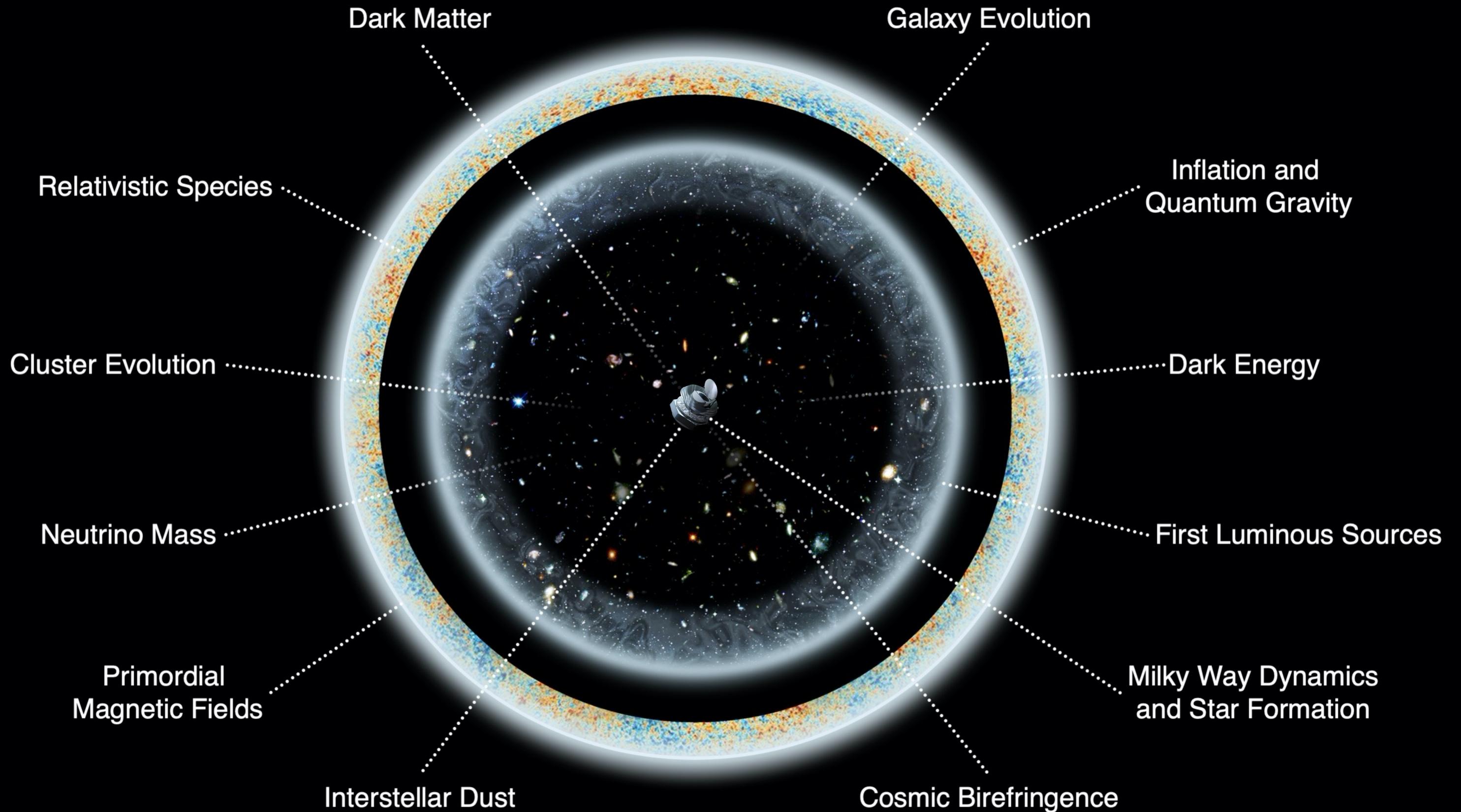
PICO echnologies are based on extensions of technologies currently used with space and sub-orbital instruments

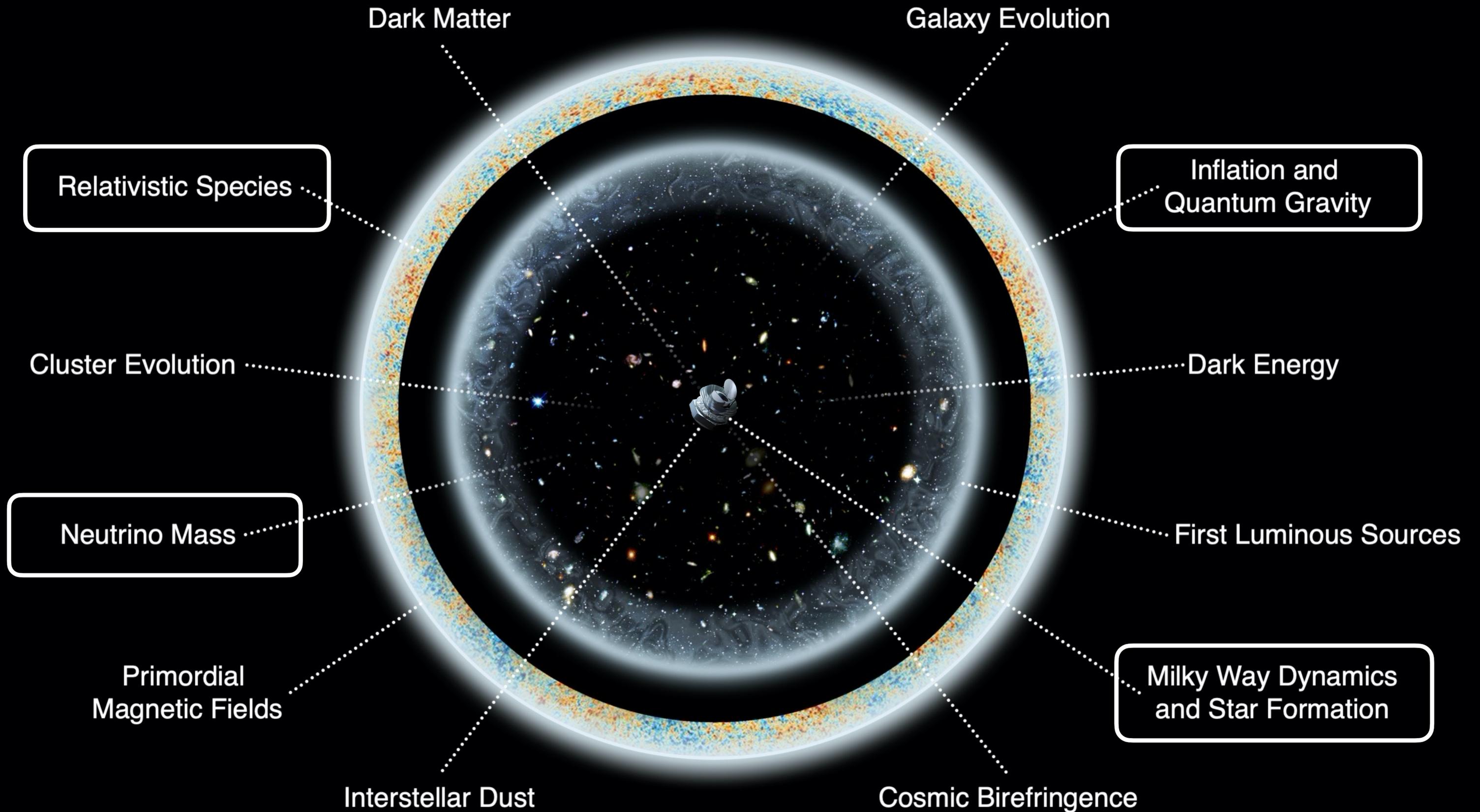


Sun Direction

Figure: JPL







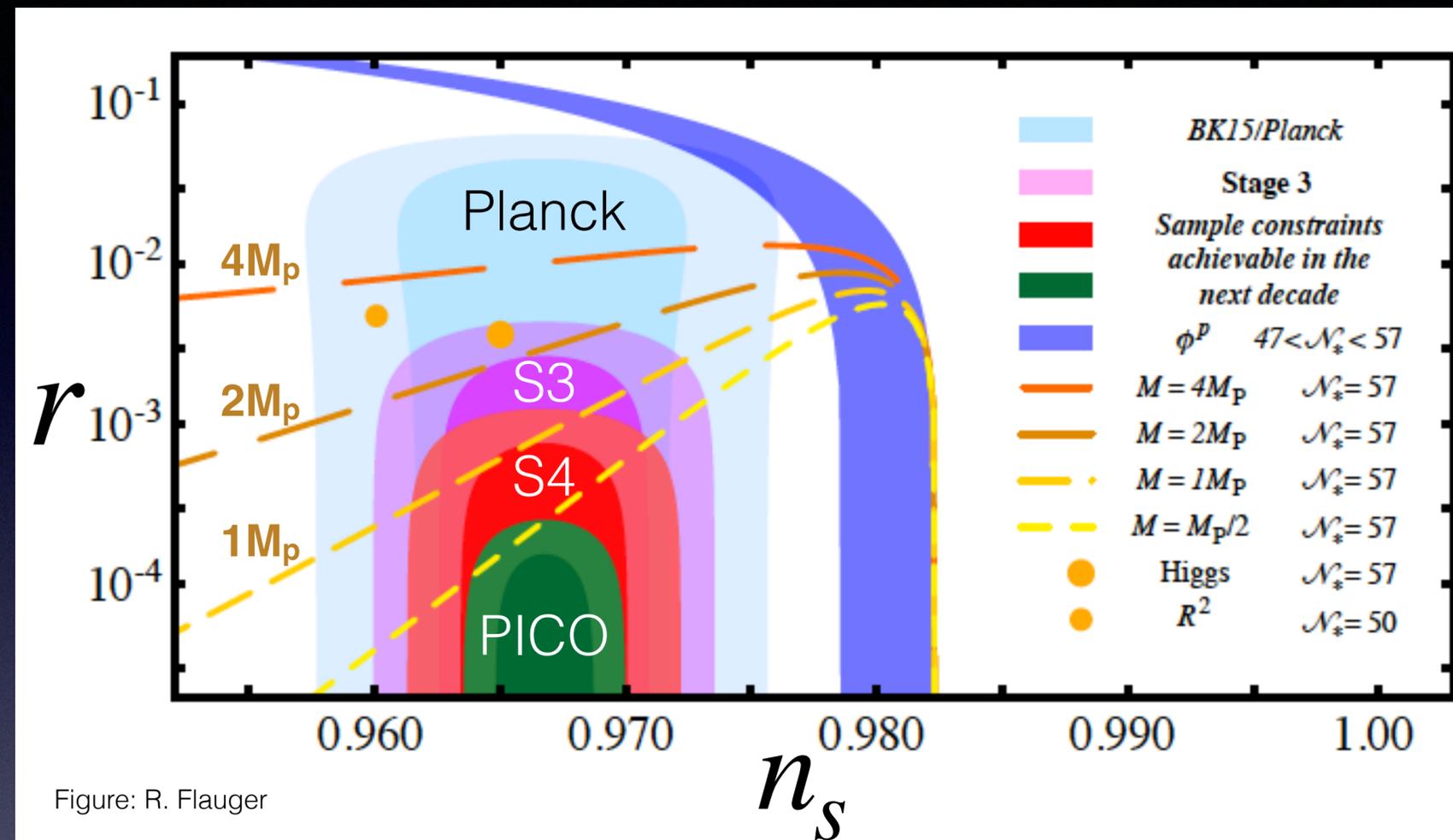
PICO S01: Tightest Constraint on Inflation r

- Textbook Inflation models that naturally explain the spectral index and have super-Planckian mass have:

$$r \gtrsim 5 \times 10^{-4}$$

- PICO requirement:

$$r < 2 \cdot 10^{-4} \text{ (95\%); } r = 5 \cdot 10^{-4} \text{ (5}\sigma\text{)}$$



Only the PICO exclusion will reject all models with superPlanckian scale in the potential with high confidence

“If this threshold is passed without detection, most textbook models of inflation will be ruled out, and the data would force a significant change in our understanding of the primordial Universe”
(Shandera et al. 2019, Community endorsed decadal white paper)

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Only the PICO exclusion will reject all models with superPlanckian scale in the potential with high confidence

- If $r \sim 1 \times 10^{-3}$ - PICO has:
 - Systematics control: Highest SNR, most stable thermal platform, simplest design
 - Foreground control: Multiple detections in independent patches of the sky

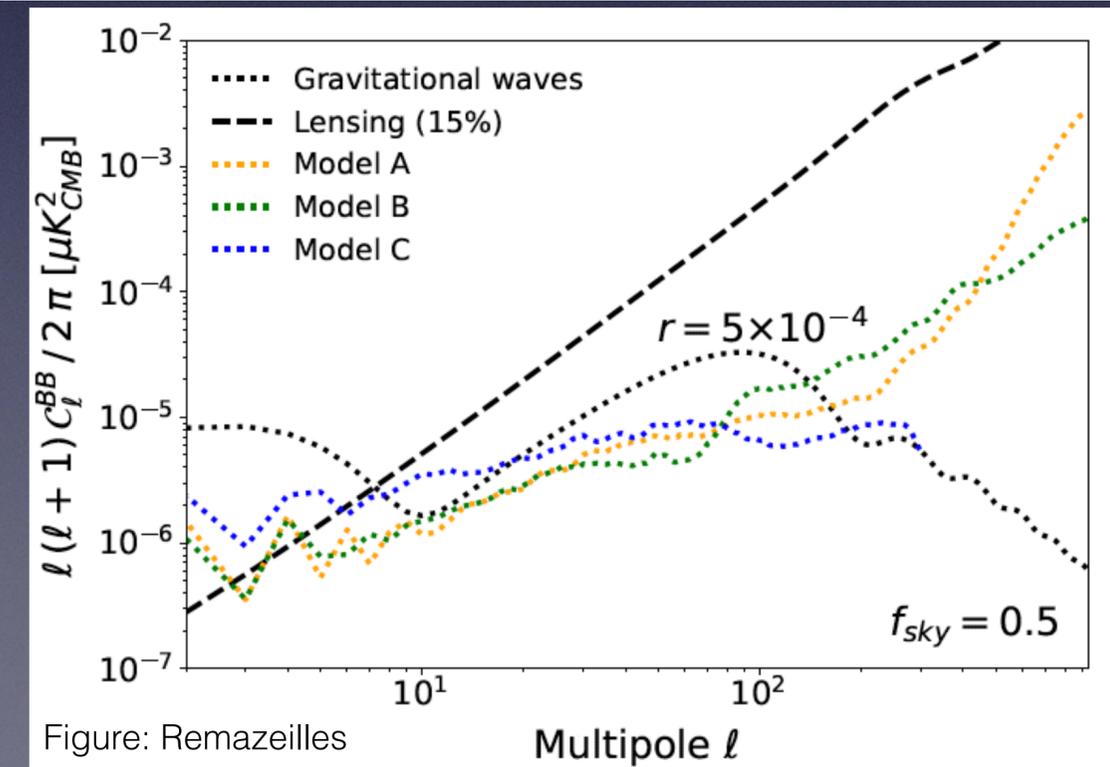
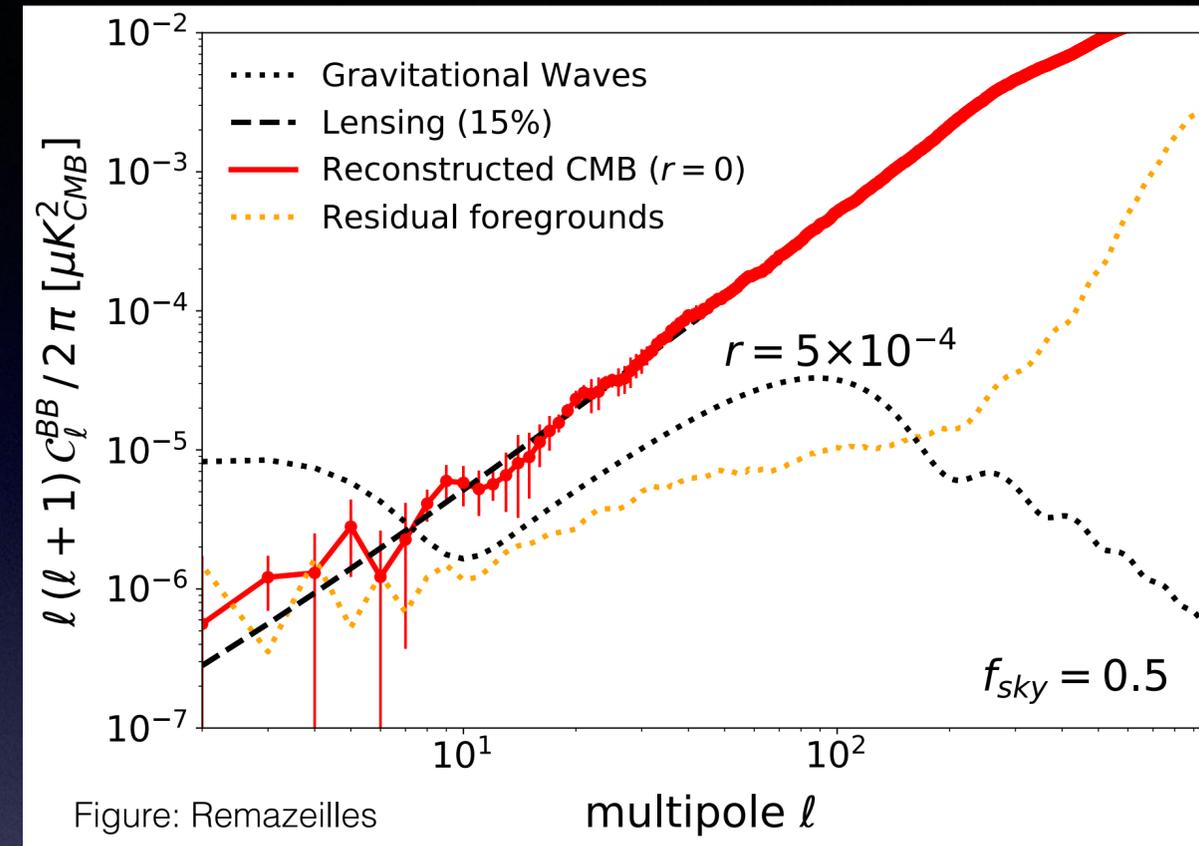
“If this threshold is passed without detection, most textbook models of inflation will be ruled out, and the data would force a significant change in our understanding of the primordial Universe”
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Can the Foregrounds be Handled

- Fisher forecast that includes correlated foregrounds, foreground separation, 40% sky, and delensing gives $\sigma(r) = 2 \times 10^{-5}$

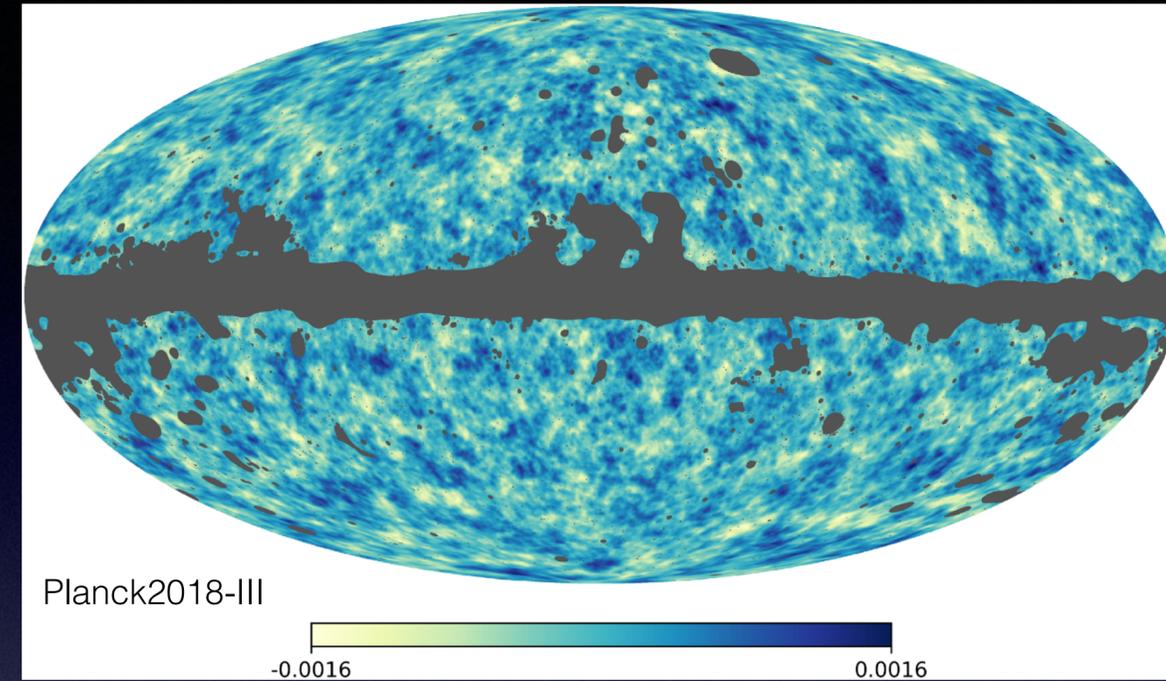
Can the Foregrounds be Handled

- Map based simulations (PySM + others), $r=0$, 50% of sky, 15% lensing, PICO noise, GNILC foreground removal with 21 bands
- Lowest ℓ has x2 bias relative to lensing, x10 lower than $r = 5 \times 10^{-4}$ (5σ)
- For $\ell=100$, residual is x4 lower
- Results approximately reproduced with other models



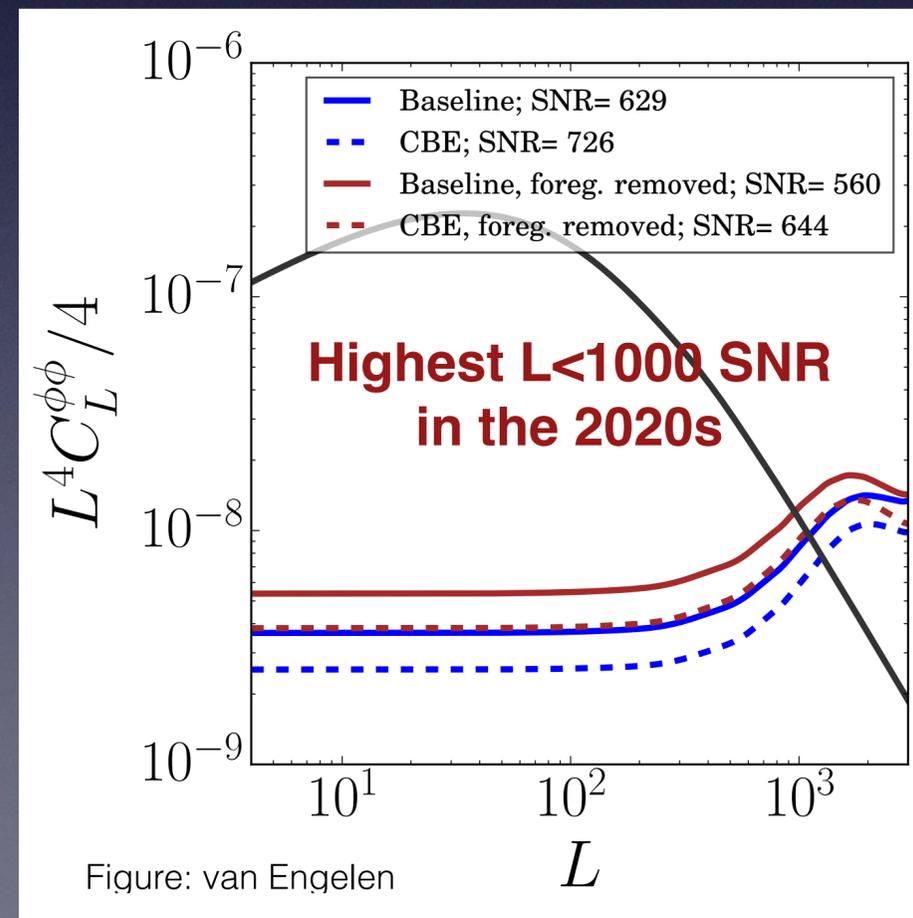
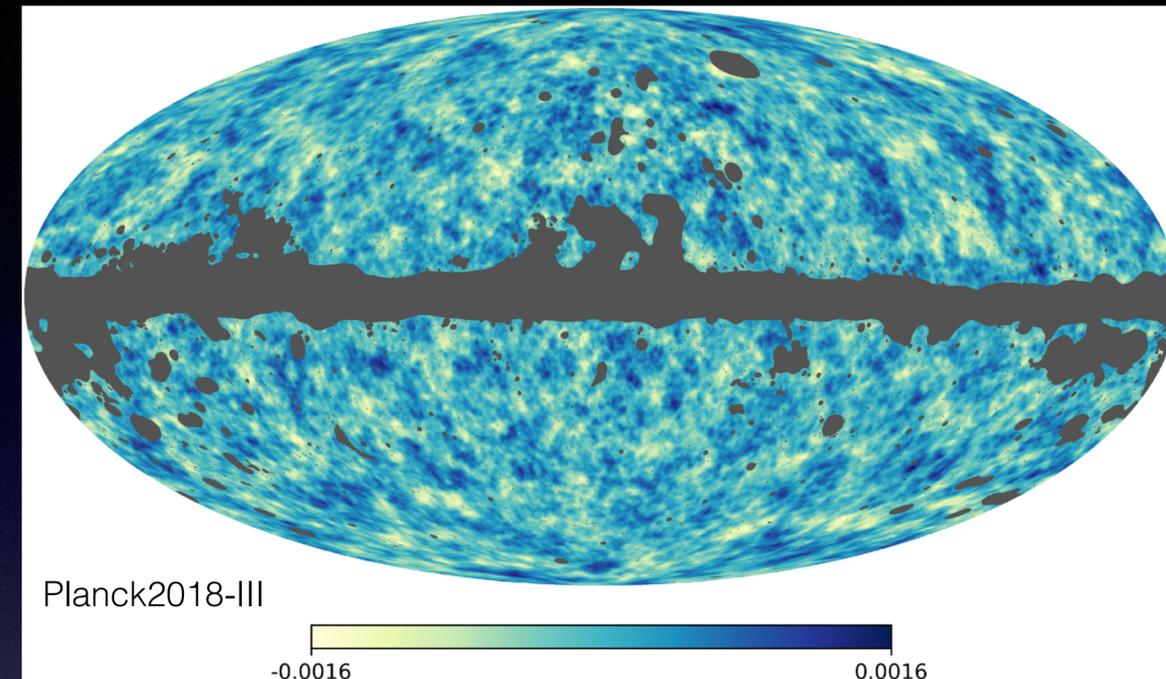
PICO S03: 4σ Detection of Neutrino Mass

- Only cosmology can determine the absolute mass scale if it is near the minimum allowed sum $\Sigma m_\nu = 58 \text{ meV}$
- Growth of structure is affected by neutrino mass, and the projected gravitational potential - revealed through CMB lensing maps - is a sensitive probe of the growth of structure



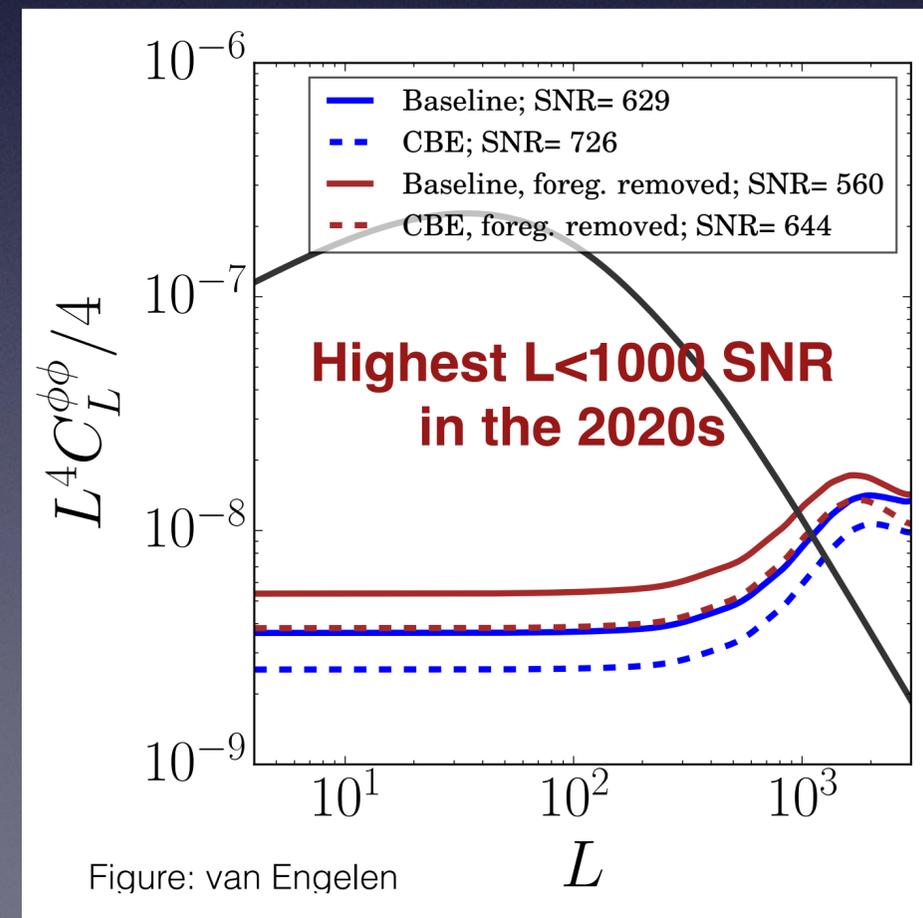
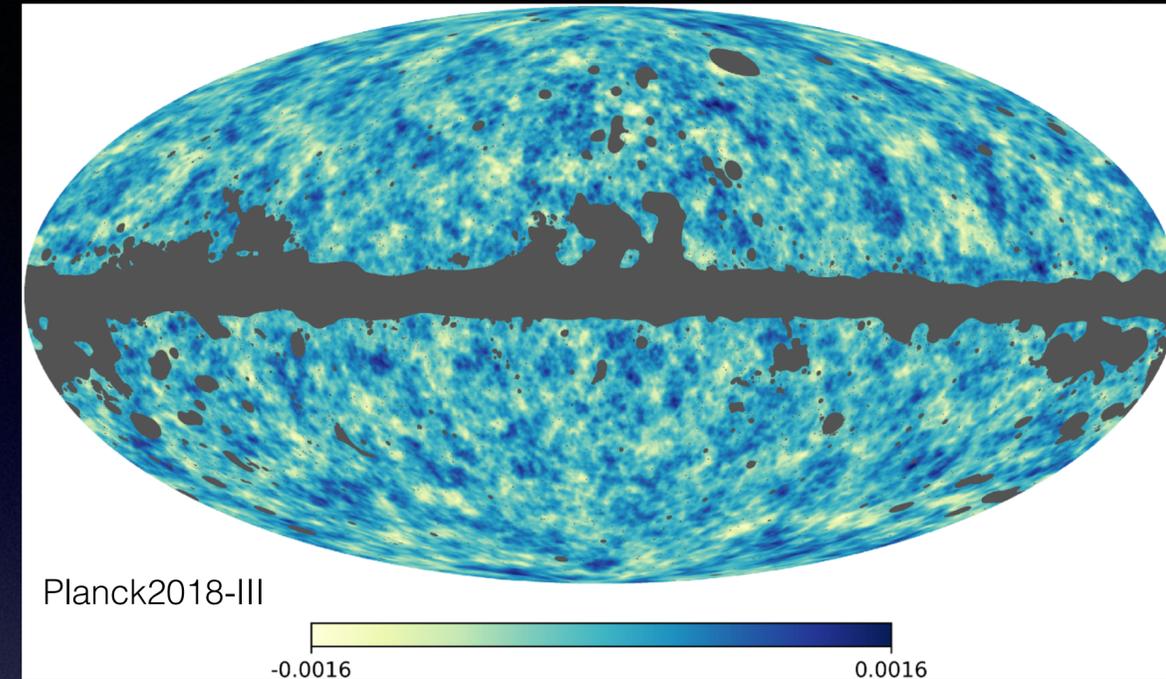
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- Growth of structure is affected by neutrino mass, and the projected gravitational potential - revealed through CMB lensing maps - is a sensitive probe of the growth of structure
- Sum of neutrino mass requires:
 - Matter density (Baryon acoustic oscillations: DESI/Euclid)
 - Growth of structure (PICO SNR=560; *Planck* SNR=40)
 - Optical depth to reionization (PICO $\sigma(\tau) = 0.002$)



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- $\sigma(\Sigma m_\nu) = 14 \text{ meV}$, ($4\sigma = 56 \text{ meV}$), one of three independent constraints



PICO SO3: 4σ Detection of Neutrino Mass

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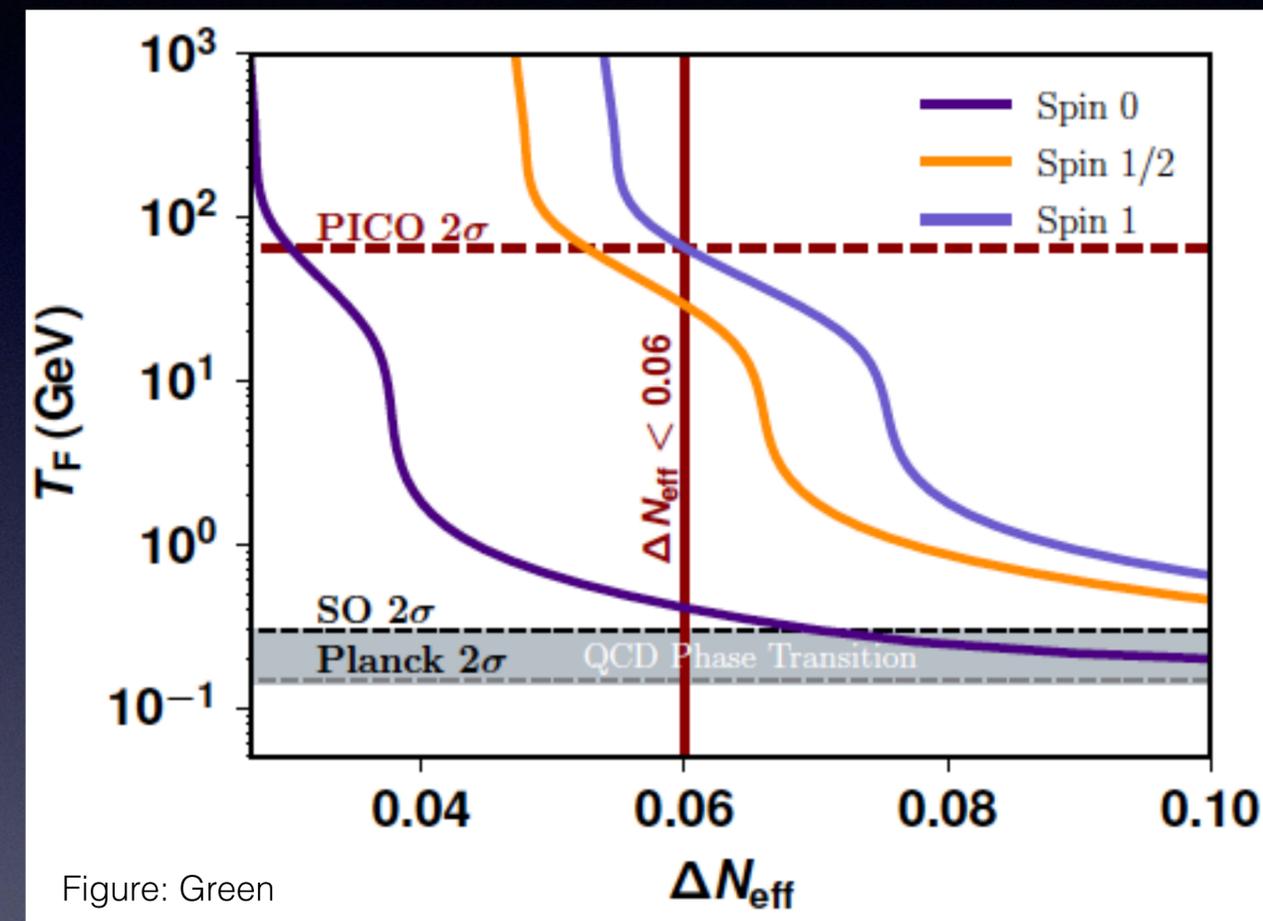
Only PICO can provide two of the three inputs within a consistent, self-calibrated dataset

No other constraint is expected to be tighter

PICO S04: New Particles

- Light species, beyond 3 neutrinos, could have existed in the early universe and fallen out of thermal equilibrium at high temperature T_F .
- CMB spectra are sensitive to the number of light species N_{eff}
- Only 3 neutrinos gives: $N_{\text{eff}} = 3.046$
- Planck + BAO: 2.92 ± 0.36 (95%)
- PICO: $\Delta(N_{\text{eff}}) = 0.06$ (95%)

Decoupling temperature as a function of ΔN_{eff} relative to neutrinos only for additional particle species



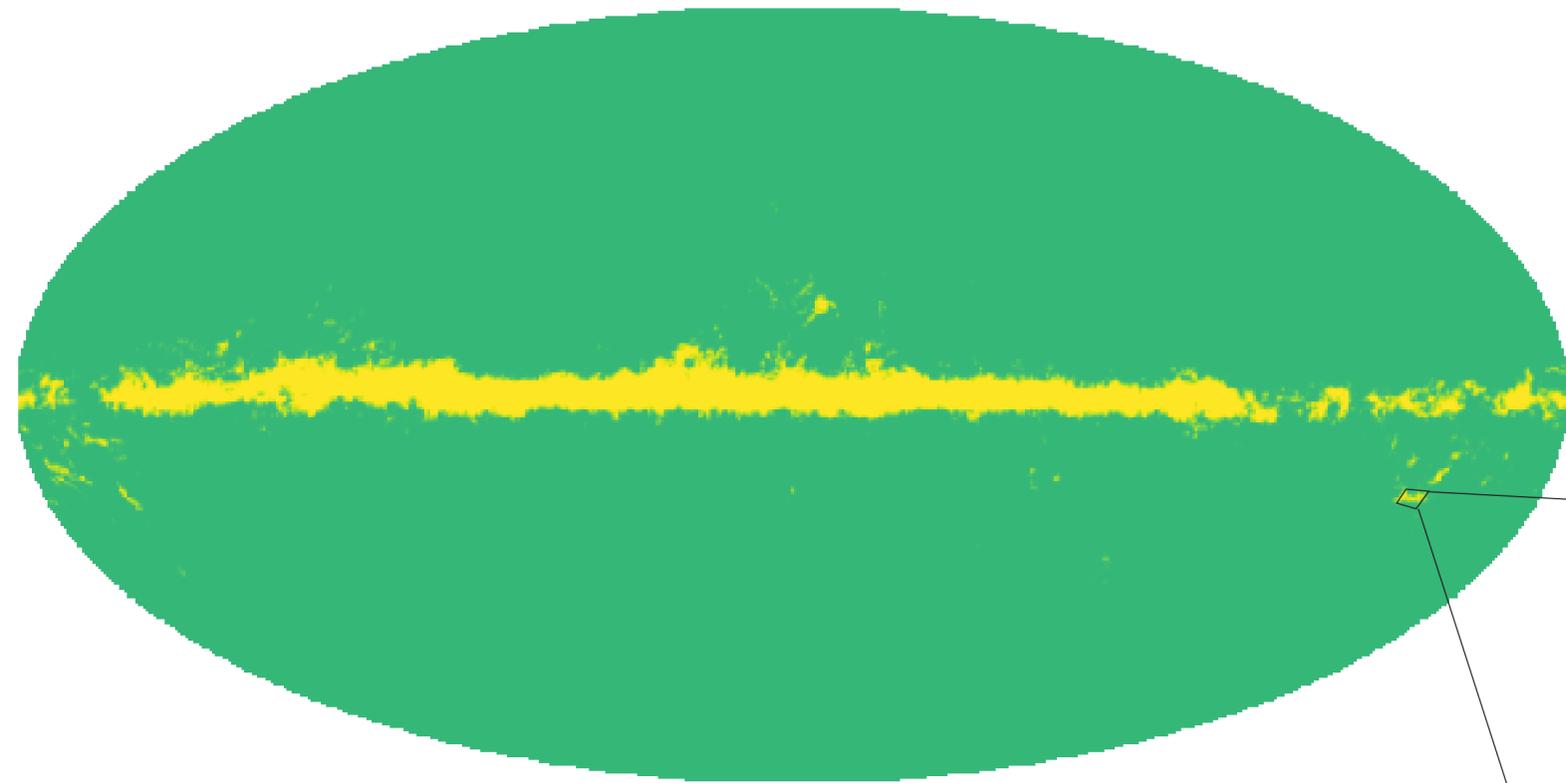
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PICO S07: Why the Low Star Formation Efficiency?

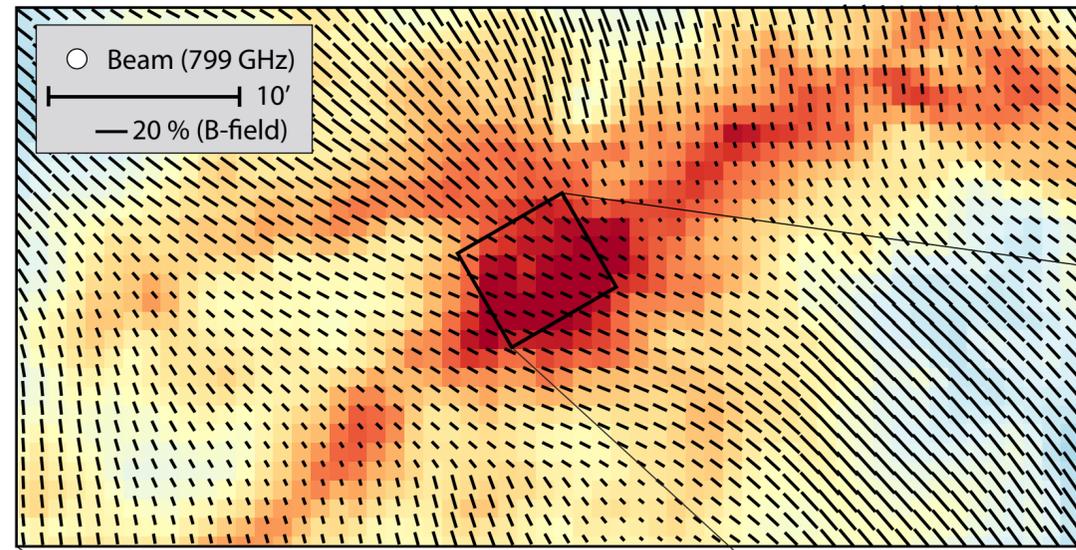
- Milky Way stars form at much lower rate than would be expected from gravitational collapse
- Turbulence + magnetic fields slow collapse from the diffuse ISM to molecular clouds, to star forming regions
- What is the ratio of energy stored in the magnetic field to that stored in turbulent motion over spatial scales from the diffuse ISM to dense cores?
- Need measurements of magnetic fields over four orders of magnitude: entire galaxy (10^4 pc) down to dense cores (0.1-1 pc)

PICO S07: Why the Low Star Formation Efficiency?

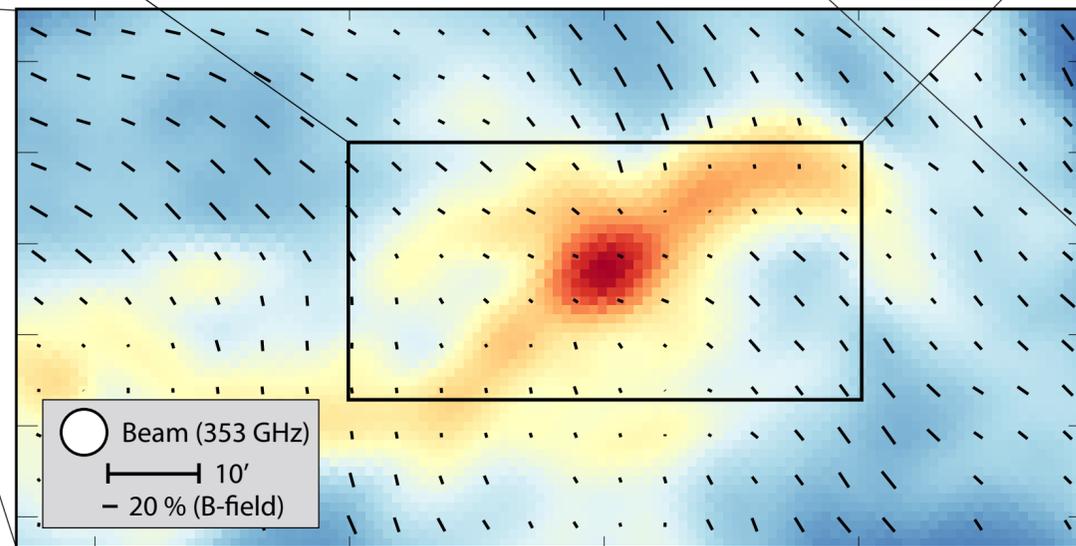
86,000,000 independent B field measurements
x1000 more than Planck



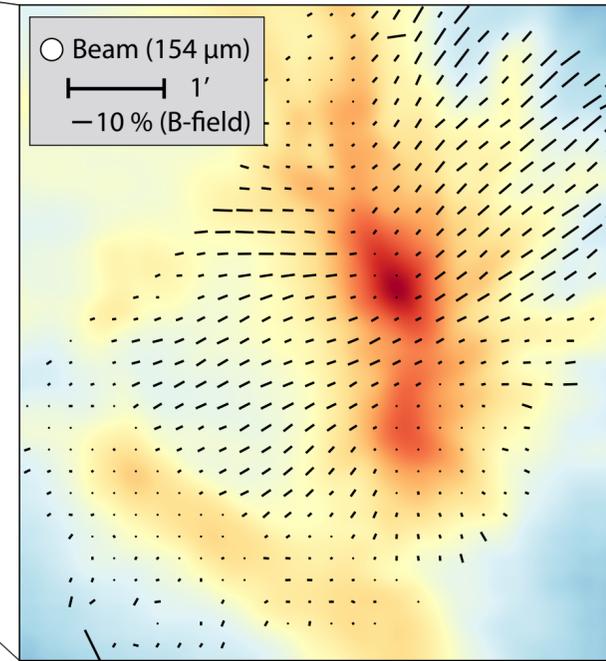
- Planck 353 GHz polarization 5' resolution, $\sigma_p < 0.67\%$
- PICO 799 GHz polarization 1' resolution, $\sigma_p < 0.67\%$



PICO (1')



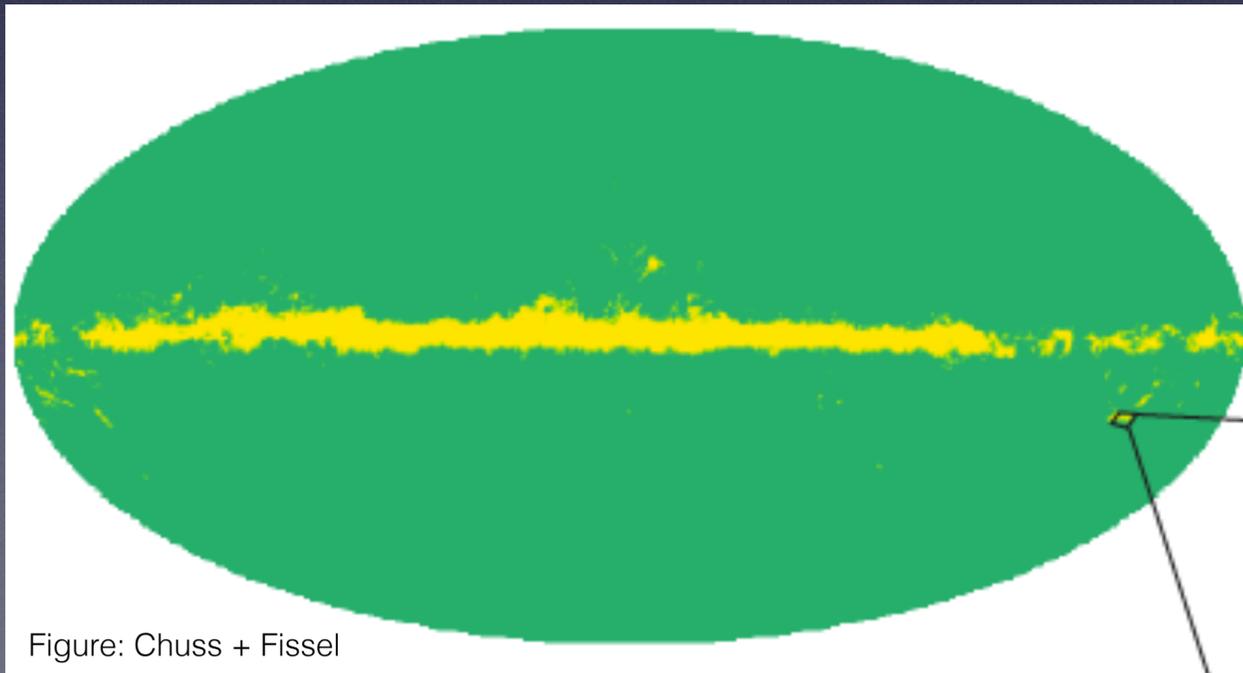
Planck (5')
Orion Region



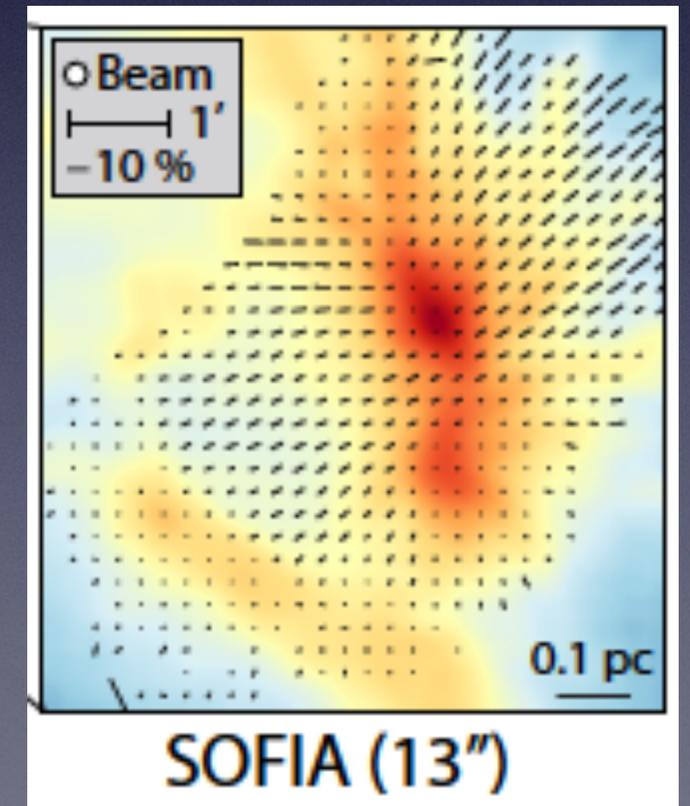
SOFIA (13'')

PICO Science : Galactic Magnetic fields

- Map magnetic fields in 70 external galaxies, with 100 measurements per galaxy (currently 2 are mapped)
- Map 10 nearby clouds with 0.1 pc resolution => scale of cloud cores (currently no data are available to connect magnetic fields in the diffuse ISM to that in cloud cores)



Factor of 10^4 in spatial scale



PICO S07: Why the Low Star Formation Efficiency?

86,000,000 independent B field measurements
x1000 more than Planck



Only PICO can generate such a dataset

- Planck 353 GHz polarization 5' resolution, $\sigma_p < 0.67\%$
- PICO 799 GHz polarization 1' resolution, $\sigma_p < 0.67\%$

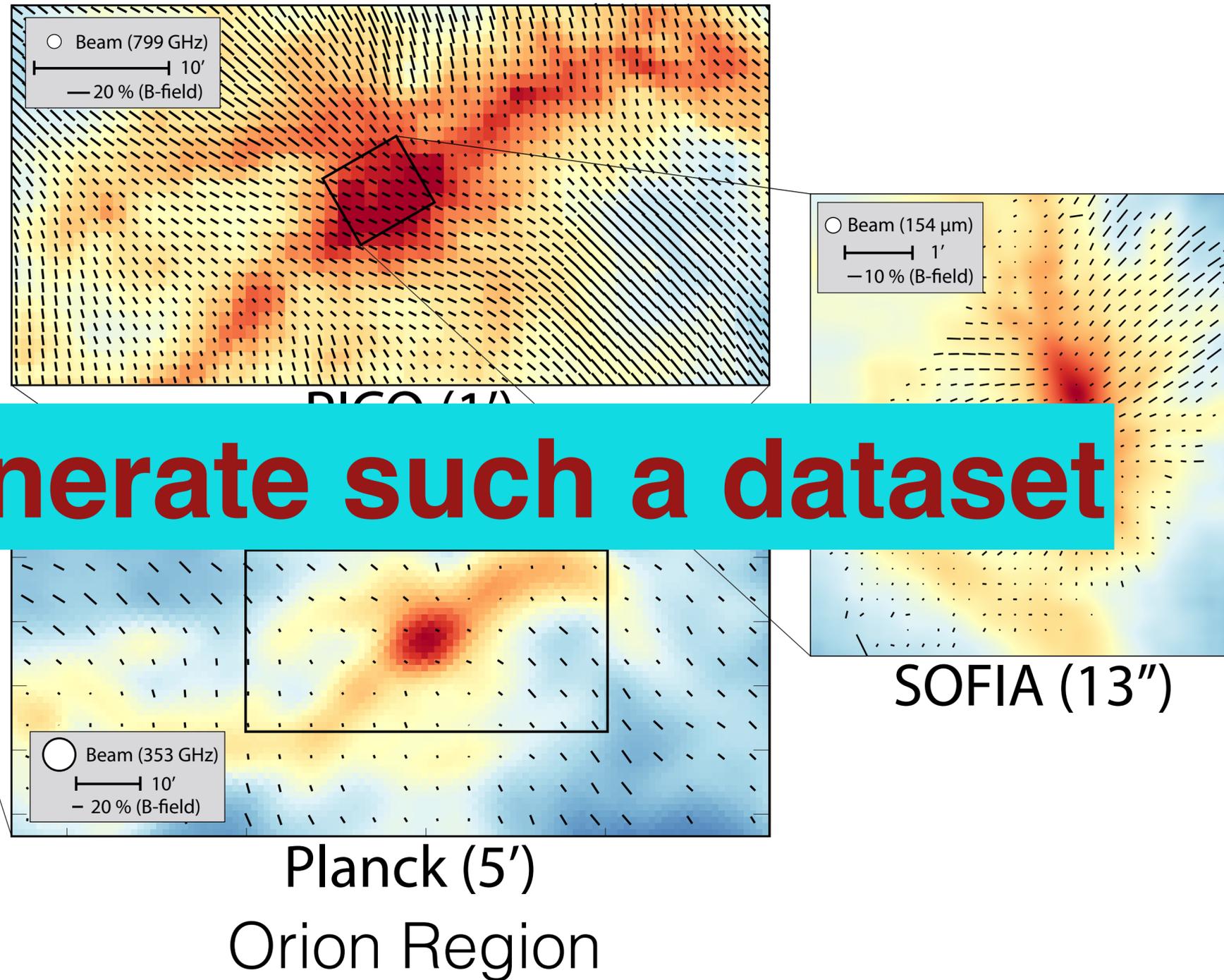
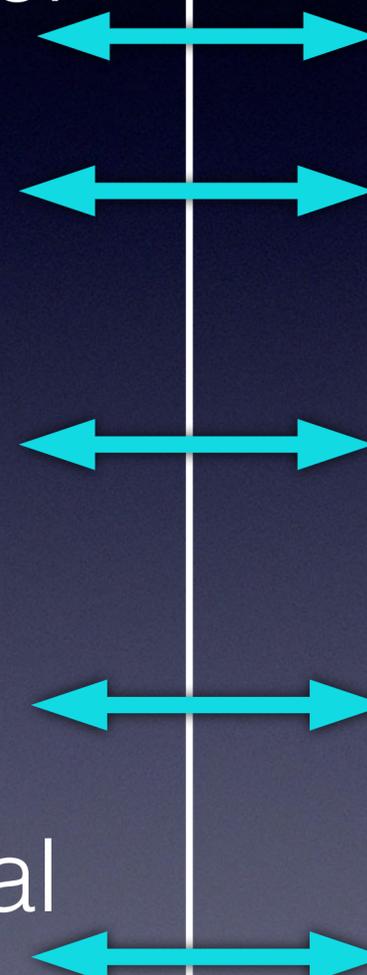


Figure: Chuss + Fissel

Legacy Surveys Available only with PICO Data

Science

- Early galaxy formation and dark matter substructure
- Early cluster formation
- Correlation of dust with galaxy properties
- Physics of jets in radio sources
- Ordering of magnetic fields in external galaxies



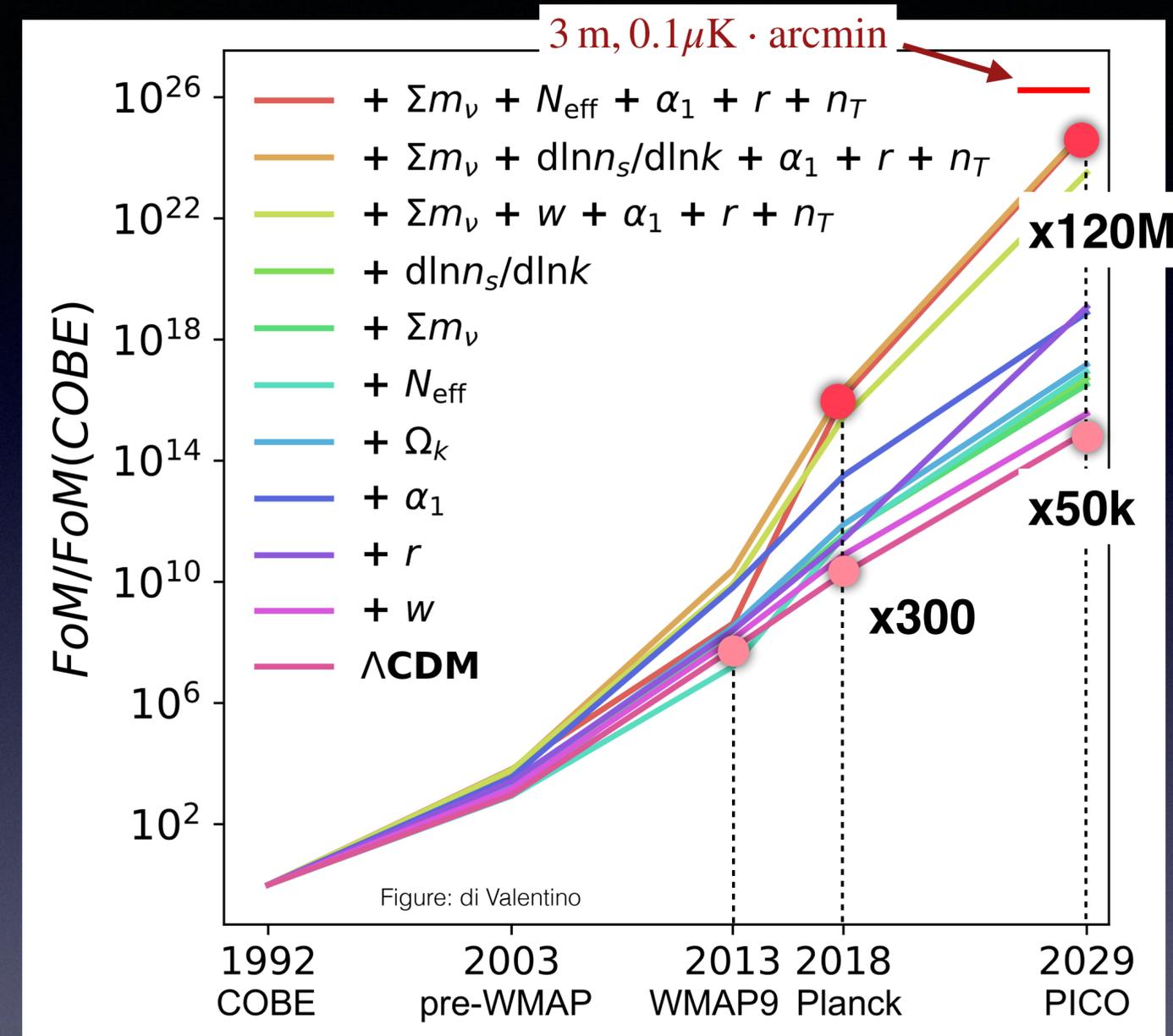
Catalog

- 4500 strongly lensed galaxies, $z \sim 5$; (x400)
- 50,000 proto-clusters, $z \sim 4.5$; (x1000)
- 30,000 galactic dust SEDs, $z < 0.1$; (x10)
- 2000 polarized radio sources; (x10)
- Polarization of few thousand dusty galaxies (x1000);

Data will be mined for years by astrophysicists in many sub-disciplines

Set Cosmological Paradigm for the 2030s

- 6-parameter Λ CDM describes the Universe well
- But tensions exist
 - 4σ between supernovae and CMB measurements of H_0
 - 2σ in measurements of σ_8 (amplitude of fluctuations)
 - What is most of the Universe made of?
- Constraint on 6-parameter Λ CDM:
 - PICO/Planck = 50,000 (Planck/WMAP9 = 300)
- Constraint on 11-parameter Λ CDM+:
 - PICO/Planck = 1.2×10^8



Λ CDM will either survive this stringent scrutiny, or a new cosmological paradigm will emerge

PICO's Status

- 50 pg PICO report publicly available (astroph/1902.10541)
- Project paper submitted in 7/2019 (astroph/1908.07495)
- Additional information has been provided to the sub-panel on Electromagnetic Observations from Space II (12/2019)
- Work on PICO will be restarting (initial focus is on foregrounds)

Why PICO, Why Now

- Transformative science; Much of the science can only be done from space.
- Further progress with CMB requires a leap in sensitivity, foreground characterization, and systematic control. Space is best suited to provide this leap.
- PICO is the only instrument with the combination of sky coverage, resolution, frequency bands, and sensitivity to achieve all of the science with one platform.
- Next decade ground-based efforts are equivalent in cost to PICO. With more bands, higher sensitivity, better control of systematics, and simpler instrument implementation, PICO is the most cost effective path for progress.

Figure: R. Flauger

Dark Matter

Evolution

213 Authors and Endorsers

Quantum Gravity

Dark Energy

Relativistic Species

Cluster Evolution



Endorsers

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Mustafa A. Amin	Jeff Booth	Ken Ganga	Ely Kovetz	Joseph Mohr	Giampaolo Pisano	Anže Slosar	Licia Verde
Adam Anderson	Sean Bryan	Tuhin Ghosh	Kerstin Kunze	Lorenzo Moncelsi	Nicolas Ponthieu	Tarun Souradeep	Patricio Vielva
James Annis	Carlo Burigana	Sunil Golwala	Guilaine Lagache	Pavel Motloch	Giuseppe Puglisi	Suzanne Staggs	Abigail Vieregg
Jason Austermann	Giovanni Cabass	Riccardo Gualtieri	Daniel Lenz	Tony Mroczkowski	Benjamin Racine	George Stein	Jan Vrtilek
Carlo Baccigalupi	Robert Caldwell	Jon E. Gudmundsson	François Levrier	Suvodip Mukherjee	Christian Reichardt	Radek Stompor	Benjamin Wallisch
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Marco Bersanelli	Chang Feng	Marc Kamionkowski	Tomotake Matsumura				
Federico Bianchini	Ivan Soares Ferreira	Reijo Keskitalo	Darragh McCarthy				
Daniel Bilbao-Ahedo	Aurelien Fraisse	Rishi Khatri	P. Daniel Meerburg				

Interstellar Dust

Cosmic Birefringence

Extra Slides